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**A CCD LIGHT CURVE OF THE ECLIPSING
CATACLYSMIC VARIABLE PG0027+260**

Ultraviolet extra object PG0027+260(B=14^m8) was reported by Green et al. (1982) as cataclysmic variable candidate. Light curve evidence is necessary to confirm it.

The CCD light curve of PG0027+260 was observed in R color on 30 November, 1989 with the 60/90/180 cm Schmidt telescope at Beijing Observatory, using the BAO CCD light curve survey photometer system(Wei Mingzhi et al. 1989). We chose two stars within the CCD field as comparison and check stars respectively, denoted by C1 and C2. Figure 1 shows the light curves of check star minus comparison and variable minus comparison. Each point in the light curves is a one-hundred-second integration and the interval between each integration is 12 seconds.

The light curve shows definitely that PG0027+260 is an eclipsing cataclysmic binary system. The depth of primary minima is about 0^m5. Like other cataclysmic variable system, the rapid changes in the brightness, which is considered to be caused by mass transfer from the secondary, can be seen in the light curve. The amplitude of rapid changes reach to 0^m2. The secondary minima is too weak to be distinguished from the strong rapid variations. This object is the second eclipsing cataclysmic variable confirmed by the BAO Schmidt telescope CCD light curve survey photometer system, the other one is PG0818+513(Chen Jiansheng et al. 1989).

Two times of primary minima were obtained, they are:

JD2447861.0689

JD2447861.2154

The orbital period is about 0^d146. We gave the ephemeris as follows,

$$\text{Min.I.} = \text{JD2447861.0689} + 0^{\text{d}}146\text{E}$$

The further observation and detailed study on physical model of this object is in process.

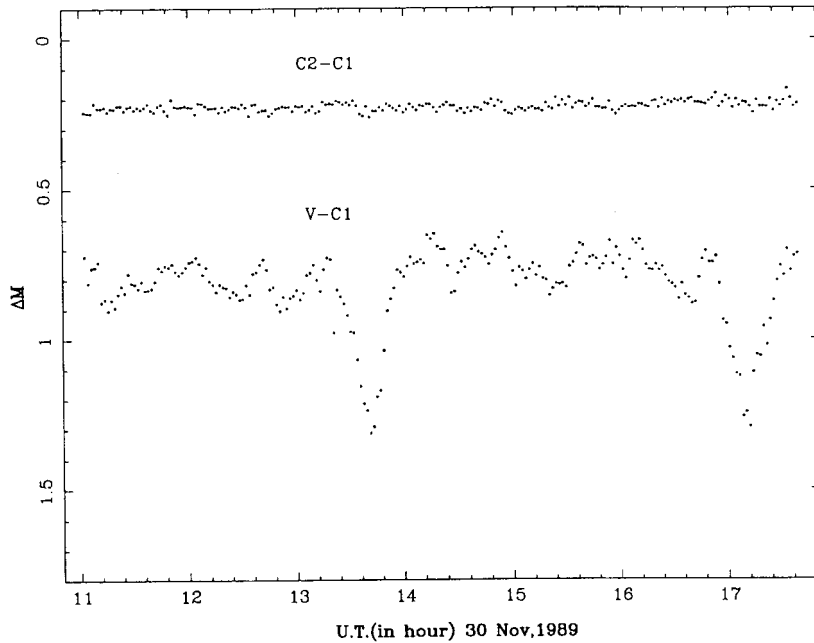


Figure 1

CCD light curves in R color, upper is light curve of C2-C1, lower is that of V-C1. Each point is a one-hundred-second integration, the interval between each integration is 12 seconds. The uncertainty of each point is $0^m.01$.

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